



# Olivine-rich asteroids in the near-Earth space

Marcel Popescu,<sup>1,2★</sup> D. Perna,<sup>1,3</sup> M. A. Barucci,<sup>1</sup> S. Fornasier,<sup>1</sup> A. Doressoundiram,<sup>1</sup>  
C. Lantz,<sup>1,4</sup> F. Merlin,<sup>1</sup> I. N. Belskaya<sup>1,5</sup> and M. Fulchignoni<sup>1</sup>

<sup>1</sup>LESIA, Observatoire de Paris, PSL Research University, CNRS, Univ. Paris Diderot, Sorbonne Paris

Cité, UPMC Univ., Paris 06, Sorbonne Universités, 5 Place J. Janssen, Meudon Pricipal Cedex F-92195, France

<sup>2</sup>Astronomical Institute of the Romanian Academy, 5 Cuștil de Argint, 040557 Bucharest, Romania

<sup>3</sup>INAF-Osservatorio Astronomico di Roma, Via Frascati 33, I-00078 Monte Porzio Catone, Rome, Italy

<sup>4</sup>IAS, UMR 8617, CNRS, Université Paris-Sud, Bat 121, F-91405 Orsay, France

<sup>5</sup>Institute of Astronomy, V.N. Karazin Kharkiv National University, 35 Sum'ska Str, UA-61022 Kharkiv, Ukraine

Accepted 2018 March 5. Received 2018 February 14; in original form 2017 December 13

## ABSTRACT

In the framework of a 30-night spectroscopic survey of small near-Earth asteroids (NEAs), we present new results regarding the identification of olivine-rich objects. The following NEAs were classified as A-type using visible spectra obtained with 3.6-m New Technology Telescope: (293726) 2007 RQ17, (444584) 2006 UK, 2012 NP, 2014 YS34, 2015 HB117, 2015 LH, 2015 TB179, 2015 TW144. We determined a relative abundance of 5.4 per cent (8 out of 147 observed targets) A-types at a 100-m size range of NEA population. The ratio is at least five times larger compared with the previously known A-types, which represent less than  $\sim 1$  per cent of NEAs taxonomically classified. By taking into account that part of our targets may not be confirmed as olivine-rich asteroids by their near-infrared spectra, or they can have a nebular origin, our result provides an upper-limit estimation of mantle fragments at size ranges below 300 m. Our findings are compared with the ‘battered-to-bits’ scenario, claiming that at small sizes the olivine-rich objects should be more abundant when compared with basaltic and iron ones.

**Key words:** methods: observational – techniques: spectroscopic – minor planets, asteroids: general.

## 1 INTRODUCTION

The asteroids with olivine-rich compositions provide fundamental clues for understanding the accretion and geochemical evolution of primitive bodies (e.g. de León et al. 2004). They are expected to be formed (e.g. Sanchez et al. 2014) either through magmatic differentiation, being the major constituent of the mantles of most differentiated bodies (e.g. Burbine, Meibom & Binzel 1996), or through nebular processes which can produce olivine-dominated objects like the R-chondrite parent body (Schulze et al. 1994; Sunshine et al. 2007).

In the optical region, the reflectance spectra of olivine-rich asteroids show very steep to extremely steep slope at wavelengths short-ward than  $0.75\ \mu\text{m}$  and a moderately deep absorption feature long-ward (Tholen & Barucci 1989; Bus & Binzel 2002a). They are classified as A-type (e.g. Veeder, Matson & Tedesco 1983; Bus & Binzel 2002b; DeMeo et al. 2009) which is an end-member taxonomic class.

It was found that these objects have moderately high albedos and a spectral band centred around  $1.05\ \mu\text{m}$  (Sunshine et al. 2007) which is composed of several absorptions generated by the presence of  $\text{Fe}^{2+}$  in the olivine structure (Burns 1993). The  $2\text{-}\mu\text{m}$  feature characteristic for pyroxene compositions is absent or very weak which suggests almost a pure olivine composition on the surface (Sanchez et al. 2014). Olivine exists also in the S-types asteroid population (associated with olivine–pyroxene assemblages). But, it is a difference between the S-type olivine asteroids and olivine-dominated A-type asteroids which reflect real geologic differences in the parent body history and the formation processes (Gaffey et al. 2015).

A very few olivine-rich asteroids (about 30 according to various data sets), which could be a representative of the mantle of disrupted differentiated bodies, are known to exist (e.g. DeMeo et al. 2015). This observational result is unexpected when compared with the meteoritic evidences which show that at least 100 chondritic parent bodies in the Main Belt experienced partial or complete melting and differentiation before being disrupted (Scheinberg et al. 2015; Scott et al. 2015, and references therein) – thus they should have produced a much larger number of olivine-dominated objects. The paucity of these objects in the Main Belt is called ‘Missing Mantle Problem’

\* E-mail: [Marcel.Popescu@obspm.fr](mailto:Marcel.Popescu@obspm.fr), [popescu.marcel1983@gmail.com](mailto:popescu.marcel1983@gmail.com)

(DeMeo et al. 2015). Several scenarios have been proposed as a solution: (i) the ‘battered-to-bits’ scenario (Burbine et al. 1996) proposes that the mantle and crustal material of the original differentiated bodies had been ground down to pieces below the limit of detectability of current spectroscopic surveys; (ii) Elkins-Tanton, Weiss & Zuber (2011) proposed that the classic view of asteroids differentiating into a pyroxene-rich crust, olivine-rich mantle, and iron core may be uncommon. They suggest that only the parent body core heats and melts due to the  $^{26}\text{Al}$  decay, but the exterior remains unheated and primitive and it is hiding the evidence of interior differentiation. In addition to these models, there is the problem of multiple scattering of the objects relative to the place where they formed during the early history of Solar system formation (DeMeo et al. 2015).

A total of 17 asteroids were identified as A-types during the SMASSII survey (Bus & Binzel 2002b). In order to define the limits of this class, they used the five asteroids unambiguously identified by Tholen taxonomy as A-types, namely (246) Aspöck, (289) Nenetta, (446) Aeternitas, (863) Benkoela, and (2501) Lohja. The SMASSII spectrum of (289) Nenetta shows a subtle absorption feature centred near  $0.63\ \mu\text{m}$  which may be attributed to various phenomena without an obvious relation to the composition (Sunshine & Pieters 1998).

About 12 near-infrared spectra of A-type asteroids are analysed by Sanchez et al. (2014). This sample includes some of the previous objects classified based on visible spectral data. They distinguished two classes, one called monomineralic olivine which exhibits only a  $1\text{-}\mu\text{m}$  feature and the second for which the spectra exhibit the  $1\text{-}\mu\text{m}$  feature and a weak  $2\text{-}\mu\text{m}$  feature corresponding to pyroxene. They did not find any link of the type core-mantel-crust between these objects and asteroids families.

To confirm the A-type candidates from Sloan Digital Sky Survey (SDSS) Moving Object Catalog, DeMeo et al. (2014) conducted a near-infrared spectral survey of asteroids. They report another 20 A-type asteroids (F. DeMeo, personal communication, February 2018) throughout the main asteroid belt, tripling the number of known A-types. They found that A-types represent only 0.16 per cent of the Main Belt population larger than  $\sim 1\ \text{km}$  in diameter and they estimated that there are about 600 of these objects with  $H \leq 17$ , half of which in the absolute magnitude range 16–17. New large spectro-photometric surveys in the near-infrared region (Popescu et al. 2016) allow us to identify new A-type candidates at these size ranges.

Using taxonomic classification of more than 1000 near-Earth asteroids (NEAs) from various sources, Binzel, Reddy & Dunn (2015) found that the A-types are uncommon in this population, representing a fraction less than 1 per cent. They also noted that A- and E-types are more common among Mars-crossers ( $\approx 5$  per cent) which can indicate a slow diffusion into Mars-crossing orbits from the Flora or Hungaria regions.

We present new spectroscopic results regarding the identification of olivine-rich asteroids in the small NEA population (absolute magnitudes higher than 20). Due to the fact that these objects are faint (typically their visual apparent magnitudes are fainter than 18), the accessible way to recognize them is by means of visible spectroscopy. The spectra discussed in this paper were obtained during a 30-night optical spectroscopic survey of NEAs, as part of the NEOShield-2 project (Perna et al. 2018). The results of the observing programme allow us to estimate the ratio of A-type NEAs smaller than 300 m.

The article is organized as follows. Section 2 describes the observing programme and the data reduction performed in order to

obtain the spectra. Section 3 introduces the methods used to analyse the asteroid spectra. The probability of associating these asteroids with olivine-dominated compositions is discussed. Section 4 shows the results with respect to the physical and dynamical parameters of each object. The implications of our findings are shown in Section 5. The ‘Conclusions’ section ends the article.

## 2 OBSERVATIONS AND DATA REDUCTION

The observations reported in this article were made in the framework of NEOShield-2 project.<sup>1</sup> The main objectives of the corresponding work-package were to undertake an extensive observational campaign involving complementary techniques with the aim to provide physical and compositional characterization of a large number of near-Earth objects (NEOs) in the 100-m size range. We performed a spectroscopic survey at ESO 3.6-m New Technology Telescope to obtain information about the composition of the small population of asteroids (Perna et al. 2018). The survey was divided in 12 observing sessions spread almost uniformly (one session every 2–3 months) over 2 yr (2015 April–2017 February). Each session consisted in two or three observing nights. This format allowed us to cover the majority of asteroids with an estimated diameter lower than 300 m and brighter than  $V \approx 20.5$  (where  $V$  is the apparent magnitude) at the observing date.

The EFOSC2 instrument was used with the Grism #1 diffraction element. A slit width of 2 arcsec was selected considering an average seeing of 1–1.5 arcsec. This configuration covers the spectral interval  $0.4\text{--}0.92\ \mu\text{m}$  with a resolution of  $R \approx 500$ . All the observations were made by orienting the slit along the parallactic angle, to minimize the effects of atmospheric differential refraction. The strategy was to observe all asteroids as close to the zenith as possible. The G2V solar analogues were observed in the apparent vicinity and at similar airmasses with the objects of interest. The number of solar analogues observed each night was larger than the number of asteroids with the purpose to avoid any artefacts caused by G2V selection and observation.

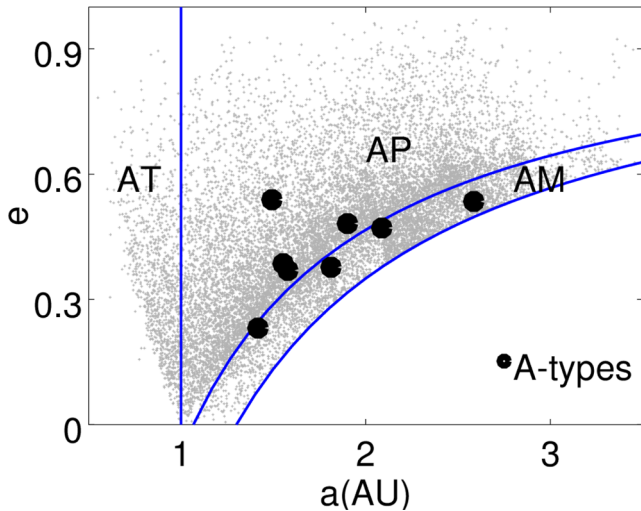
The data reduction followed the standard procedures (bias and background subtraction, flat-field correction, one-dimensional spectra extraction). We used OCTAVE software to generate IRAF scripts for performing these tasks automatically. The APALL package was used to extract the one-dimensional spectra across the apertures. Wavelength calibration was obtained using emission lines from the He-Argon lamp. In the final step, the reflectivity of the objects was determined by dividing the asteroid spectra with those of the solar analogues. For each asteroid, we selected the solar analogue which gives a result close to the median when comparing with the results obtained with respect to all other solar analogues observed during that night. All spectra were normalized at  $0.55\ \mu\text{m}$  and the results are presented in Fig. 2.

Spectral observations are subject to multiple selection biases due to asteroid properties (size, albedo, and distance to observer), and due to equipment (Jedicke, Larsen & Spahr 2002). Due to the incompleteness of observed versus as-yet undiscovered asteroids at this size range (our survey doubled the number of 100-m size range asteroids which can be taxonomically classified), any de-biasing technique required to derive statistics between groups with very different properties (e.g. albedos) is poorly constrained. However, because we did a random selection of objects, those with similar sizes and albedos had equal probabilities to be observed.

<sup>1</sup> <http://www.neoshield.eu/>

**Table 1.** Observing conditions: asteroid designation, the orbit type (AM: Amor; AP: Apollo; \*: potentially hazardous asteroid), semimajor axis, eccentricity, inclination, minimum orbital intersection distance with the Earth (MOID), mid UTC of the observations, the apparent magnitude ( $V$ ), the phase angle ( $\alpha$ ), the airmass, the total exposure time (Exp.), and the corresponding solar analogue (S.A.) used to get the reflectance, are presented.

Designation	Orbit	$a$ (au)	$e$	$i$ ( $^\circ$ )	MOID (au)	UTC	$V$ (mag)	$\alpha$ ( $^\circ$ )	Airmass	Exp. (s)	S.A.
(293726) 2007 RQ17	AP	1.579	0.37	2.0	0.021	2015-06-08T07:44	18.1	51	1.288	$2 \times 600$	SA 107998
(444584) 2006 UK	AP*	1.493	0.54	4.7	0.014	2016-05-12T06:28	16.5	29	1.731	$1 \times 600$	SA 107998
2012 NP	AM	2.088	0.47	7.4	0.098	2015-07-20T02:17	18.9	41	1.356	$2 \times 900$	HD 111244
2014 YS34	AP*	1.553	0.39	6.9	0.007	2015-06-09T07:14	20.0	20	1.249	$2 \times 1125$	SA 1021081
2015 HB117	AM	1.417	0.23	6.5	0.081	2015-06-08T04:24	20.2	33	1.359	$2 \times 1200$	SA 1021081
2015 LH	AP	1.902	0.48	3.9	0.003	2015-06-09T06:30	19.1	44	1.426	$1 \times 180$	SA 1021081
2015 TB179	AM	2.586	0.53	8.1	0.226	2015-11-05T07:34	20.0	27	1.203	$3 \times 1200$	HD 11123
2015 TW144	AM	1.813	0.38	8.4	0.201	2015-11-05T05:46	19.3	22	1.285	$2 \times 900$	Hyades64



**Figure 1.** The observed A-type asteroids in a ( $a$ ,  $e$ ) representation of NEAs ( $a$  is the semimajor axis and  $e$  the eccentricity). The blue lines are the limits of different orbital classes (AM-Amor, AP-Apollo, AT-Aten).

In this article, we focus on the asteroids classified as A-type. This was the case of eight objects, namely: (293726) 2007 RQ17, (444584) 2006 UK, 2012 NP, 2014 YS34, 2015 HB117, 2015 LH, 2015 TB179, and 2015 TW144. Their orbital elements and the observation details are provided in Table 1. Fig. 1 shows that this sample has no particular orbital elements with respect to the known population of NEAs. In addition, the orbital inclination,  $i$ , of these asteroids is in the  $2$ – $8.5^\circ$  range. The asteroids (444584) 2006 UK and 2014 YS34 are catalogued by Minor Planet Center (MPC) as a potentially hazardous asteroid (PHA).

### 3 METHODS USED TO ANALYSE DATA

The A-type is a distinct end-member class in most of the recent taxonomies (e.g. Tholen 1984; Barucci et al. 1987; Tholen & Barucci 1989; Fulchignoni, Birlan & Antonietta Barucci 2000; Bus & Binzel 2002b; DeMeo et al. 2009). Based on the spectra obtained over  $0.44$ – $0.92 \mu\text{m}$  spectral interval, Bus & Binzel (2002b) provide the following definition: ‘Very steep to extremely steep UV slope short-ward of  $0.75 \mu\text{m}$ , and a moderately deep absorption feature, long-ward of  $0.75 \mu\text{m}$ . A subtle absorption feature is often present around  $0.63 \mu\text{m}$ ’.

Several authors (e.g. Gaffey et al. 2015) argue that some of the A-types identified based on visible spectra do not show near-infrared data compatible with olivine-dominated asteroids. We note that from 17 objects reported in SMASII survey (Bus & Binzel 2002a)

**Table 2.** The two spectral classes from Bus taxonomy that best matched our spectra and their mean square differences (msq).

Designation	Tax 1	msq.tax1	Tax 2	msq.tax2
(293726) 2007 RQ17	A	0.0109	Ld	0.0153
(444584) 2006 UK	A	0.0007	Sa	0.0012
2012 NP	A	0.0024	Sa	0.0049
2014 YS34	A	0.0008	Sa	0.0024
2015 HB117	A	0.0015	Ld	0.0024
2015 LH	A	0.0060	Sa	0.0102
2015 TB179	A	0.0007	Sa	0.0015
2015 TW144	A	0.0016	Sa	0.0030

as belonging to A-type asteroids, 13 have near-infrared data up to  $2.45 \mu\text{m}$ . In this sample, six objects are classified as A-type: (246) Asporina, (289) Nenetta, (446) Aeternitas, (863) Benkoela, (1951) Lick, and (2501) Lohja (de León et al. 2004; DeMeo et al. 2009). The asteroid (4142) Dersu-Uzala is reported as A-type by Binzel et al. (2004a) but displays a  $2\text{-}\mu\text{m}$  feature which indicates a non-negligible pyroxene content. Different classifications are provided for (3352) McAuliffe – Binzel et al. (2004b) reported it as A-type, and Whiteley (2001) as S-type. The asteroids (1126) Otero, (2715) Mielikki, (4713) Steel, and (5641) McCleese are classified as Sw-type and (2732) Witt as L-type by DeMeo et al. (2009). Although this statistics is extremely poor, it provides a rough estimation of the accuracy for recognizing olivine-dominated assemblages based on visible data.

Within the limits allowed by our data (spectral interval and signal-to-noise ratio), we used several methods to ensure the correct classification as A-type for the asteroids shown in this paper: (i) we computed the mean square differences with respect to all taxonomic types defined by Bus & Binzel (2002b); (ii) we determined the position of the maximum, the slopes in the  $(0.45\text{--}0.65) \mu\text{m}$  (denoted as  $BR_{\text{slope}}$ ) and  $(0.75\text{--}0.92) \mu\text{m}$  (denoted as  $IZ_{\text{slope}}$ ) spectral intervals (Table 3) and we compared them with those of standard taxonomic types; (iii) we compared these spectral parameters with those of known A-types; and (iv) we compared our data with all spectra available in the RELAB data base (NASA Reflectance Experiment Laboratory<sup>2</sup>). These methods are presented below.

(i) The first two spectral classes from Bus taxonomy that matched our spectra according to the mean square differences are presented in Table 2 and plotted in Fig. B1. All objects are classified as A-type considering a calculated distance at half compared with the second match which is Sa-type for the majority. This is fully in agreement with an olivine-dominated spectrum.

<sup>2</sup> <http://www.planetary.brown.edu/relabdocs/relab.htm>

**Table 3.** Some characteristics of the NEAs studied in this article. The asteroid designations, absolute magnitude, estimated diameter, and the spectral slopes [ $BR_{\text{slope}}$  (per cent/0.1  $\mu\text{m}$ ),  $IZ_{\text{slope}}$  (per cent/0.1  $\mu\text{m}$ )] computed based on our observation.

Designation	$H$ (mag)	Diam.(m)	Max.pos( $\mu\text{m}$ )	$BR_{\text{slope}}$ (per cent/0.1 $\mu\text{m}$ )	$Err_{BR_{\text{slope}}}$	$IZ_{\text{slope}}$ (per cent/0.1 $\mu\text{m}$ )	$Err_{IZ_{\text{slope}}}$
(293726) 2007 RQ17	22.5	96	0.762	40.2	0.5	−13.5	1.3
(444584) 2006 UK	20.2	277	0.775	25.8	1.0	−13.4	3.5
2012 NP	21.3	167	0.763	29.8	0.5	−14.9	2.0
2014 YS34	20.8	210	0.775	25.0	1.0	−10.0	5.1
2015 HB117	23.6	58	0.762	25.8	1.0	−6.0	3.0
2015 LH	27.3	11	0.736	37.9	1.7	−21.1	3.8
2015 TB179	20.6	231	0.748	24.7	0.8	−12.7	2.3
2015 TW144	20.7	220	0.738	26.3	0.6	−12.3	1.3

(ii) The position of the peak of our identified spectra (Table 3) ranges from 0.736  $\mu\text{m}$  (which belongs to 2015 LH) to 0.775  $\mu\text{m}$  (which belongs to 2014 YS34) with an average value  $\lambda_{\text{peak}} = 0.754 \pm 0.017 \mu\text{m}$ . The  $BR_{\text{slope}}$  computed over the interval (0.45–0.65)  $\mu\text{m}$  is for the majority larger than 25 per cent/0.1  $\mu\text{m}$ . These values compare only with the A-type objects that were classified by SMASS and which have the average  $BR_{\text{slope}}^{\text{SMASS}} = 25.1 \pm 3.5$  per cent/0.1  $\mu\text{m}$ . For a comparison, the S-complex objects reported by DeMeo et al. (2009) have  $BR_{\text{slope}}^{\text{Scomp}} = 14.3 \pm 3.8$  per cent/0.1  $\mu\text{m}$ .

The presence of the band at 1  $\mu\text{m}$  is critical to differentiate between the A-types and other possible red-type spectra corresponding to L-, Ld-, and D-types from Bus taxonomy. The computed values for the  $IZ_{\text{slopes}} < -6$  per cent/0.1  $\mu\text{m}$  strongly indicate the presence of the feature even for low signal-to-noise ratio spectra. All these parameters point to an A-type classification, well distinctive from L-/Ld-type or from the most common S-type (Fig. B1).

(iii) Depending on the shape of the maximum, Bus & Binzel (2002b) identified two spectral curves associated with the A-class. The first one, which is similar with the spectrum of (289) Nenetta and (863) Benkoela, has the maximum at about  $\approx 0.75 \mu\text{m}$  sharply peaked with 1  $\mu\text{m}$  band particularly rounded while the second has a broader reflectance like the spectrum of (246) Asporina and no upward curvature out to 0.92  $\mu\text{m}$ . To assess a comparison, we computed the mean square error differences in the (0.70, 0.80)  $\mu\text{m}$  spectral interval between our asteroids spectra and the SMASSII spectra (Bus & Binzel 2002a) of (289) Nenetta, (863) Benkoela, and (246) Asporina. Within the limit of noise, we found that (293726) 2007 RQ17 and 2015 LH have a more (863) Benkoela-like spectrum and (444584) 2006 UK, 2012 NP, 2014 YS34, 2015 HB117, 2015 TB179, and 2015 TW144 have a bowl shape maximum similar with (246) Asporina. However, this is a rough comparison as most of the spectra show intermediate shapes between these references.

(iv) The comparison with RELAB spectral data base considering only the visible spectra provides few constraints with respect to the composition of the object as multiple different solutions may be found to match the same visible data. Also, it has to be noted that space-weathering effects change the spectral curve of asteroids (Brunetto et al. 2006). However, the comparison deserves to be performed as a measure of the various solutions which can be found. The most relevant results of this comparison are shown in Table A1 and plotted in Fig. A1. The comparison was performed using M4AST<sup>3</sup> tool (Popescu, Birlean & Nedelcu 2012). The solutions show that these spectral data match the different mixtures of olivine or thin sections of olivine from the meteorites. In particular, one of the RELAB spectra that fit our observed data was of a

particulate olivine fayalite mineral. Nevertheless, among spectral curve matches, there are also several corresponding to typically ordinary chondrite meteorites (e.g. L4 ordinary chondrite meteorite Barratta).

$$D = \frac{1329}{\sqrt{p_v}} 10^{-0.2H}. \quad (1)$$

The size of the asteroids presented in this paper was estimated by computing their equivalent diameter using the absolute magnitude  $H$  reported on the IAU Minor Planet Center website<sup>4</sup> and the average albedo for A-types  $p_v = 0.19$  (Mainzer et al. 2011) using the equation (1) provided by Pravec & Harris 2007. The values are reported in Table 3.

## 4 RESULTS

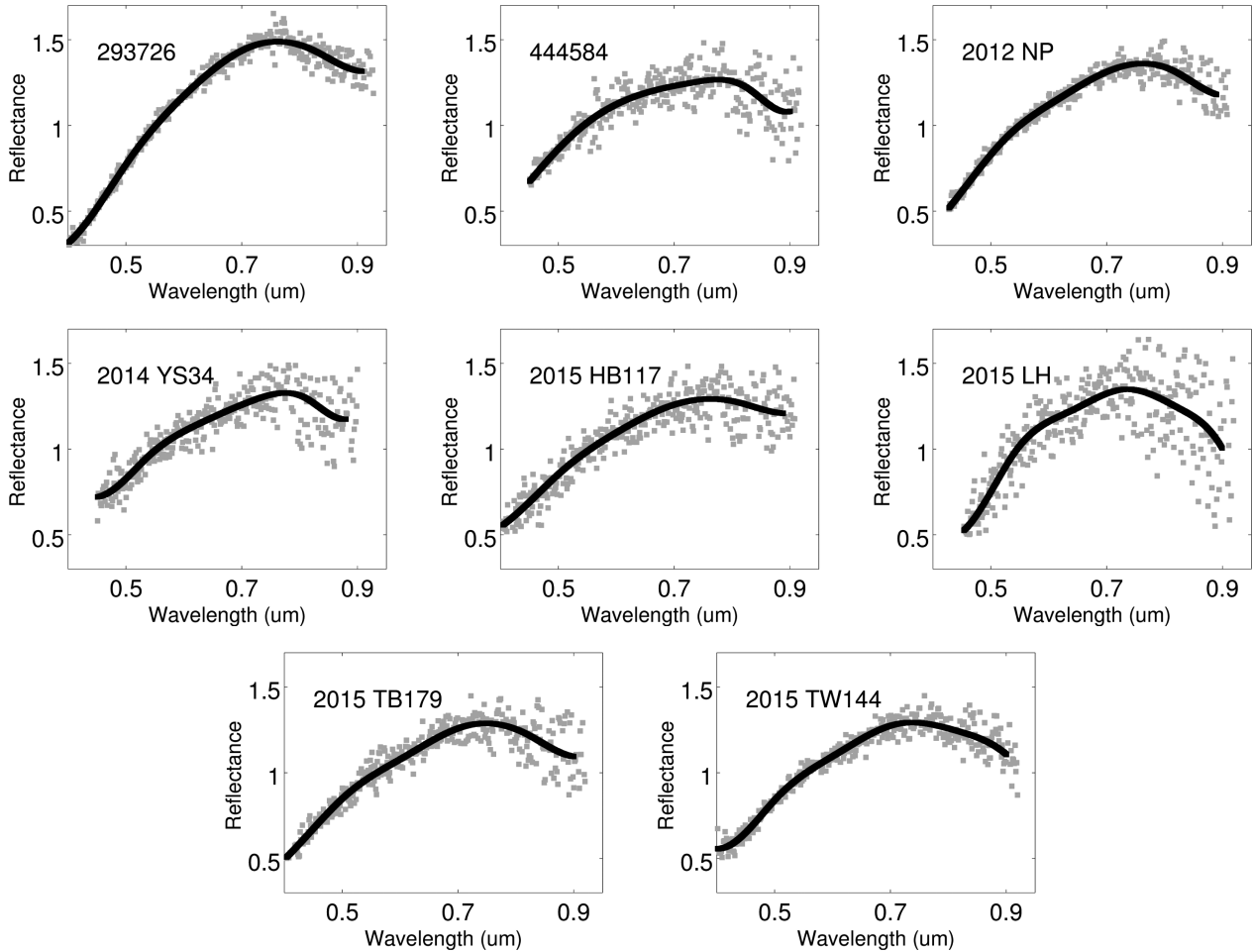
From the total sample of 147 observed NEAs with an absolute magnitude  $H \leq 20$  mag, we found eight asteroids classified as A-type. Their spectra are plotted in Fig. 2. Three of these objects have an estimated diameter less than 100 m – (293726) 2007 RQ17, 2015 HB117, and 2015 LH while the other five are  $\sim 200$  m size. Table 3 summarizes the diameters and the computed spectral parameters. In the public data bases, there is no other physical information available, except the orbital parameters and the absolute magnitude derived from astrometric data (from Minor Planet Center website). Part of these objects have sufficient astrometric data for follow up at future Earth close approaches. But the smallest objects, such as 2015 LH, were observed over a very short orbital arc and their orbit uncertainty is high. Thus, they require astrometric recovering (e.g. Vaduvescu et al. 2018).

(293726) 2007 RQ17 is an Apollo-type asteroid for which we estimated a diameter of 96 m. It was discovered in 2007 September, and although it is a 100-m size object, it has sufficient astrometric observations to receive a permanent designation (i.e. a number). It has a perihelion at  $q = 0.996$  au and the MOID (minimum orbital intersection distance) with the Earth of 0.021 au. This asteroid has the reddest spectrum among our sample – characterized by the largest  $BR_{\text{slope}}$  of about 40.2 per cent/0.1  $\mu\text{m}$ . Our observations were performed at a phase angle of  $\alpha = 51^\circ$ , thus a phase reddening factor may account for this high slope (Perna et al. 2018). No spectral feature can be distinguished below 0.7  $\mu\text{m}$ . The mineral spectrum matching the observed data is the one of an olivine fayalite sample.

(444584) 2006 UK is a PHA with an estimated diameter of 277 m and it is the largest object from the sample reported here. It has an MOID with the Earth of 0.014 au and it is on an eccentric orbit ( $e = 0.54$ ), with a perihelion  $q = 0.689$  au. It has a broad

<sup>3</sup> <http://m4ast.imcce.fr/>

<sup>4</sup> <http://www.minorplanetcenter.net>



**Figure 2.** Visible spectra of (293726) 2007 RQ17, (444584) 2006 UK, 2012 NP, 2014 YS34, 2015 HB117, 2015 LH, 2015 TB179, 2015 TW144. All spectra are normalized to 0.55  $\mu\text{m}$ . A spline-fit of the data is shown with black while the values extracted from the images are shown in grey.

maximum at 0.775  $\mu\text{m}$  and its spectral curve is the closest to the A-type template. The spectrum shows a change of slope around 0.65  $\mu\text{m}$ , which may suggest a possible feature at these wavelengths.

*2012 NP* is an Amor-type object with an equivalent diameter of 167 m. It was observed at  $\alpha = 41^\circ$  and it has a  $BR_{\text{slope}}$  of about 29.8 per cent/0.1  $\mu\text{m}$ , thus being the second reddest object. This is in agreement with the correlation between  $BR_{\text{slope}}$  and phase angle for the A-types, outlined by Perna et al. (2018).

*2014 YS34* is an Apollo-type asteroid with an equivalent diameter of 210 m. It is catalogued as a PHA. The spectrum shows a slight change of slope around 0.65  $\mu\text{m}$  which can indicate a possible feature. For this object MPC records astrometric positions covering an arc length over 888 d, thus allowing follow-up observations.

*2015 HB117* is an Amor-type object with an estimated diameter of 58 m. Astrometric positions are reported over 83 d. The feature at 0.9  $\mu\text{m}$  is less pronounced compared with the other objects which is quantified by the small value  $IZ_{\text{slope}} = -6 \pm 3$  per cent/0.1  $\mu\text{m}$ . Due to this fact, a classification as an L-type cannot be totally excluded (being at the limit of noise).

*2015 LH* is the smallest object from the observed sample, with an estimated diameter of 11 m. Its observational arc covers only 5 d. Our spectroscopic observations were performed when this asteroid was at 0.011 au geocentric distance. Although this spectrum has the lowest signal-to-noise ratio ( $\approx 11$ ), its  $BR_{\text{slope}}$  is fully compatible with an A-type classification. It shows a change of slope around

0.65  $\mu\text{m}$ . The shape of the spline-fit indicates a feature at these wavelengths.

*2015 TB179* is a 231-m diameter Amor-type object. It has the largest semimajor axis  $a = 2.586$  au and its aphelion  $q = 3.968$  au in the outer part of the Main Belt. The spectrum is well consistent with the mean A-type spectrum.

*2015 TW144* is a 220-m Amor-type object. Its spectrum shows a broad peak at 0.738  $\mu\text{m}$  and a weak feature around 0.45  $\mu\text{m}$ .

## 5 DISCUSSIONS

Within the NEA population only two asteroids have been recognized as olivine-dominated NEAs (Sanchez et al. 2014), namely (1951) Lick (de León et al. 2004) – a 5.57-km object which belongs to a Mars-crossing asteroid population, but its orbital elements (i.e. perihelion distance  $q = 1.305$  au) are close to borderline of NEAs (i.e.  $q \leq 1.3$ ); and (136617) 1994 CC (Reddy et al. 2011) which is a triplet NEA with a diameter of 650 m for the primary, 10 m for the secondary, and 5 m for the tertiary component (IAUC 9053).

By using visible and near-infrared data, de León et al. (2004) found for (1951) Lick a surface composed by almost pure olivine, estimating the Fosterite content to be  $Fo = 90 \pm 10$  per cent (low iron content). Their result is comparable with the one reported by Sanchez et al. (2014), who found  $Fo = 70 \pm 5$  per cent. Based on near-infrared data acquired with SpeX at IRTF telescope, Reddy

et al. (2011) found that (136617) 1994 CC has an Mg-rich ( $F_0 = 90$ ) olivine composition, similar to Mg-rich pallasites with a low-metal component.

In the existing spectral data bases and those of physical properties of asteroids, there are very few other NEAs reported to have a taxonomic type compatible with an olivine-rich composition. We mention only those that are not ambiguously classified (inconsistent classification when considering data reported by different authors). Thus, based only on spectral data in the visible region (471240) 2011 BT15 (Kuroda et al. 2014) and (488515) 2001 FE90 (Hicks et al. 2009) are reported as A-types. Both objects have an equivalent diameter in the range of 200–300 m ( $H = 21.7$  mag for 2011 BT15 and  $H = 20.1$  mag for 2001 FE90). The next type which may correspond to an olivine-rich composition is Sa. The objects (136849) 1998 CS1, (275677) 2000 RS11, and 1993 TQ2 (see EARN data base<sup>5</sup> and corresponding references for a detailed description) are classified as Sa. Carry et al. (2016) report an A-type object within a sample of 230 NEAs based on SDSS data. Overall, we can summarize that from the total number of NEAs spectrally characterized up to now (more than 1000 according to Binzel et al. 2015, and about 735 in the EARN data base) less than 1 per cent have spectra compatible with an olivine-rich composition.

The ratio of A-types can be outlined relative to S-complex asteroids which have similar albedo values. The S-types, which show spectra similar with ordinary chondrite material, represent more than half of all measured NEAs (Binzel et al. 2015). Thus, it can be concluded to a fraction less than 2 per cent of olivine-dominated achondrite asteroids compared to those associated with ordinary chondrite assemblages. Moreover, the percentage of A-type NEAs is one order less when compared with basaltic asteroids which are considered to originate in the crust of differentiated bodies.

We note that with the exception of Lick which has a diameter of 5.57 km, the other A-type NEAs are in the 100-m size range. This implies a very small amount of pure olivine material in terms of volume and mass – considering an average density of  $3.73 \pm 1.40 \text{ g cm}^{-3}$  (Carry 2012).

Compared with these results, we found a fraction of  $\approx(5.4$  per cent) olivine-dominated asteroids within our observed sample of 147 objects with diameters less than 300 m. This result raises several questions: What are the reasons for finding more A-types within the small NEA population? Is there a difference when comparing the small-size objects with the larger ones? What is the origin of these small olivine-dominated objects?

Due to the fact that only visible data are used to classify them as A-type and the statistically low number of observed objects, several debates can be addressed concerning both the observational programme and the accuracy of the classification. These are discussed below:

(i) Risk of observational bias: By considering the ratio of A-type objects over the commonly S-types, which have similar albedo, the bias introduced due to the tendency to observe brightest objects can be neglected (when considering the same size). The ratio of S-complex asteroids within the NEOSShield-2 observed sample is  $\approx 54$  per cent (Perna et al. 2018) comparable with the one reported by Binzel et al. (2015). Thus, our findings imply a relative fraction of A-types over S-types for NEAs lower than 300 m of 10.3 per cent (eight A-types and 78 S-complex objects).

From the dynamical point of view, as shown in Fig. 1 and Table 1, their orbital elements are common for the near-Earth objects and there is no grouping in ( $a$ ,  $e$ ) orbital elements space. Nevertheless, we note that all objects have low inclinations ( $2\text{--}8.5^\circ$ ), which may point to an origin close to the reference plane.

(ii) Misclassification due to space-weathering and/or phase-angle effects: These effects may account for the reddening of spectra, making a common S-type with an ordinary chondrite composition to be confused with an olivine-dominated object. Available experimental and numerical modelling of the weathering effect on olivine-rich surfaces has shown an increase of visible slope together with considerable reduction of the  $1 \mu\text{m}$  absorption band (e.g. Kohout et al. 2014). Sanchez et al. (2012) showed that the increase of spectral slope caused by weathering and phase reddening is of the same value. These effects can lead to ambiguous classifications only for objects close to the boundaries between taxonomic classes (Sanchez et al. 2012). In our case, the phase reddening trend is well-seen and constitutes about  $0.5$  per cent  $\text{deg}^{-1}$  (Perna et al. 2018). The reddening effect does not influence the classification of our objects observed in the range of phase angles of  $20\text{--}51^\circ$ .

(iii) Ambiguous classification due to lack of the near-infrared data: The A-types identified based on visible spectra must be considered cautiously as their near-infrared data may refute this classification (e.g. Gaffey et al. 2015). Indeed, the A-types are confirmed by the lack of  $2 \mu\text{m}$  band. But as shown in the ‘Methods’ section, although there is a very low number of objects known, it can be inferred that  $\sim$ half of the asteroids classified as A-type are found to have an olivine-dominated composition.

(iv) Systematic errors introduced by the instrument or by data reduction pipeline: A complete description of the sample of NEAs observed NEOSShield-2 survey is presented by Perna et al. (2018). Because the slope is one of the main parameters that differentiates between the A-type and S-complex objects, we verified if it was affected by a systematic error. The fact that the average slope for the S-complex objects observed by NEOSShield-2 has a value of  $BR_{\text{slope}}^{\text{NS}} = 13.5 \pm 3$  per cent/ $0.1 \mu\text{m}$  comparable with  $BR_{\text{slope}}^{\text{Scomp}} = 14.3 \pm 3.8$  per cent/ $0.1 \mu\text{m}$  computed for the sample reported by DeMeo et al. (2009) gives confidence in our findings.

We searched for other spectral observations of the objects reported in NEOSShield-2 sample. We found nine objects with spectra reported by SMASS-NEO programme.<sup>6</sup> Within the limits of noise, the results are consistent and no systematic difference has been identified.

Olivine can be formed through different mechanisms: accretion of grains from an oxidized nebular region without significant post-accretionary heating (R chondrites), partial melting with extraction of a basaltic melt leaving an olivine-rich residue (brachinites), or crystallization of olivine from a melt to form an olivine mantle or olivine-rich layer (Sunshine & Pieters 1998). Small olivine objects should be common, at least as a result of collisions of the hypothesized large number of primordial differentiated objects (see the ‘Introduction’ section). In the ‘battered-to-bits’ scenario (Burbine et al. 1996, and references therein), the small size of these objects is considered as an explanation for the lack of observational data of olivine-dominated compositions.

Olivine spectral features are known to be greatly reduced in contrast by mixture with pyroxenes and other Fe-bearing mineral phases (Singer 1981). Hence, the identification of olivine-dominated

<sup>5</sup> <http://earn.dlr.de/nea/>

<sup>6</sup> <http://smass.mit.edu/minus.html>

A-type asteroids is expected to be easier for smaller bodies, as they more probably include more concentrated deposits of purest olivine phases.

NEAs are an unstable population, with a lifetime in the order of millions of years (e.g. Menichella, Paolicchi & Farinella 1996; Gladman et al. 1997; Bottke et al. 2002). Dynamical models predict that most of these objects come from the inner (~61 per cent), middle (~24 per cent), and outer (~8 per cent) parts of the Main Belt (Bottke et al. 2002). Due to their large eccentricities, they have a short collision lifetime (e.g. Bottke, Nolan & Greenberg 1993). Thus, one of the plausible scenarios for the origin of small A-types reported here is the violent collisions of differentiated bodies resulting in fragments of the mantle of a 100-m size range.

Nevertheless, the remnants of collisions of differentiated asteroids should be in a greater volume compared with the iron and basaltic asteroids. But, even if all our eight possible A-types (~5.4 per cent) are confirmed as an olivine-dominated composition, their percentage barely compares with those of basaltic asteroids in the NEA population, and it is much less than expected from the model of a differentiated object where mantle material is more abundant than iron core and basaltic crust.

We note that, although the eccentricity and semimajor axis of the observed A-types are very different, their inclinations are in a tight range of 2–8°. Thus, a common origin of these eight possible olivine-dominated objects cannot be totally excluded. Such a compositional grouping of A-type asteroids has been recently shown. The largest members of Eureka family of Mars Trojan asteroids are olivine-dominated objects that share a common origin (Borisov et al. 2017; Christou et al. 2017; Polishook et al. 2017). Using a numeric simulation, Polishook et al. (2017) show that Mars Trojans are more likely to be impact ejecta from Mars and that olivine-rich asteroids in the Hungaria population may have the same martian origin.

## 6 CONCLUSIONS

We found eight asteroids classified as A-type in a sample of 147 NEAs spectrally observed in the optical region (Perna et al. 2018). The A-type asteroids are associated with olivine-rich compositions. This result allows us to estimate a fraction of ~5.4 per cent of olivine-dominated NEAs for objects with sizes less than 300 m. This fraction is more than five times larger compared with the ratio of known A-type asteroids in the NEA population.

Although the existence of these objects involves violent collisions of large differentiated bodies, as considered by the ‘battered-to-bits’ scenario, this fraction is far less than expected to result from the predicted number of primordial differentiated objects. In the hypothesis that NEAs reflect the broad compositional distribution of Main Belt, this result of ~5.4 per cent gives an upper limit for the number of olivine-dominated asteroids at a 100-m size range.

The result presented in this paper follows the premise that the make-up of NEO population is heavily affected by asteroid disruption events that create many small fragments that then move swiftly towards the resonances into NEO space. There may not be evidence of older break-ups because those smaller fragments have cleared out of both the Main Belt and the NEO population long ago.

## ACKNOWLEDGEMENTS

This work is based on observations collected at the European Organization for Astronomical Research in the Southern hemisphere under ESO programme 095.C-0087. We acknowledge fi-

ancial support from the NEOSShield-2 project, funded by the European Union’s Horizon 2020 research and innovation programme (contract no. PROTEC-2-2014-640351). DP has received further funding from the Horizon 2020 programme also under the Marie Skłodowska-Curie grant agreement no. 664931. MP received partial funding from a grant of the Romanian National Authority for Scientific Research and Innovation, CNCS-UEFISCDI, project number PN-II-RU-TE-2014-4-2199. This project was also supported by the Programme National de Planétologie (PNP) of CNRS/INSU, co-funded by CNES. The article makes use of data published by the following websites Minor Planet Center, E.A.R.N-The Near-Earth Asteroids Data Base which is maintained and updated by Gerhard Hahn (gerhard.hahn@dlr.de), SMASS-Planetary Spectroscopy at MIT, and RELAB Spectral Database which contains spectra acquired with the NASA RELAB facility at Brown University. We thank Dr. Francesca DeMeo for the constructive and helpful suggestions.

## REFERENCES

- Barucci M. A., Capria M. T., Coradini A., Fulchignoni M., 1987, *Icarus*, 72, 304
- Binzel R. P., Birlan M., Bus S. J., Harris A. W., Rivkin A. S., Fornasier S., 2004a, *Planet. Space Sci.*, 52, 291
- Binzel R. P., Rivkin A. S., Stuart J. S., Harris A. W., Bus S. J., Burbine T. H., 2004b, *Icarus*, 170, 259
- Binzel R. P., Reddy V., Dunn T. L., 2015, in Patrick M., Francesca E. D., William F. B., eds, *The Near-Earth Object Population: Connections to Comets, Main-Belt Asteroids, and Meteorites*. University of Arizona Press, Tucson, p. 243
- Borisov G., Christou A., Bagnulo S., Cellino A., Kwiatkowski T., Dell’Oro A., 2017, *MNRAS*, 466, 489
- Bottke W. F., Jr, Nolan M. C., Greenberg R., 1993, In *Lunar and Planetary Inst., Twenty-fourth Lunar and Planetary Science Conference. Part 1: A-F*, p. 159
- Bottke W. F., Morbidelli A., Jedicke R., Petit J.-M., Levison H. F., Michel P., Metcalfe T. S., 2002, *Icarus*, 156, 399
- Brunetto R., Vernazza P., Marchi S., Birlan M., Fulchignoni M., Orofino V., Strazzulla G., 2006, *Icarus*, 184, 327
- Burbine T. H., Meibom A., Binzel R. P., 1996, *Meteorit. Planet. Sci.*, 31, 607
- Burns R. G., 1993, *Mineralogical Applications of Crystal Field Theory*. Cambridge Univ. Press, Cambridge
- Bus S. J., Binzel R. P., 2002a, *Icarus*, 158, 106
- Bus S. J., Binzel R. P., 2002b, *Icarus*, 158, 146
- Carry B., 2012, *Planet. Space Sci.*, 73, 98
- Carry B., Solano E., Eggl S., DeMeo F. E., 2016, *Icarus*, 268, 340
- Christou A. A., Borisov G., Dell’Oro A., Cellino A., Bagnulo S., 2017, *Icarus*, 293, 243
- de León J., Duffard R., Licandro J., Lazzaro D., 2004, *A&A*, 422, L59
- DeMeo F. E., Binzel R. P., Slivan S. M., Bus S. J., 2009, *Icarus*, 202, 160
- DeMeo F. E., Carry B., Binzel R. P., Moskovitz N., Polishook D., Burt B. J., 2014, *American Astronomical Society Meeting Abstracts #224*, p. 321.09
- DeMeo F. E., Alexander C. M. O., Walsh K. J., Chapman C. R., Binzel R. P., 2015, in Patrick M., Francesca E. D., William F. B., eds, *The Compositional Structure of the Asteroid Belt*. University of Arizona Press, Tucson, p. 13
- Elkins-Tanton L. T., Weiss B. P., Zuber M. T., 2011, *Earth Planet. Sci. Lett.*, 305, 1
- Fulchignoni M., Birlan M., Antonietta Barucci M., 2000, *Icarus*, 146, 204
- Gaffey M. J., Reddy V., Fieber-Beyer S., Cloutis E., 2015, *Icarus*, 250, 623
- Gladman B. J. et al., 1997, *Science*, 277, 197
- Hicks M., Lawrence K., Rhoades H., Somers J., McAuley A., Barajas T., 2009, *Astron. Telegram*, 2116

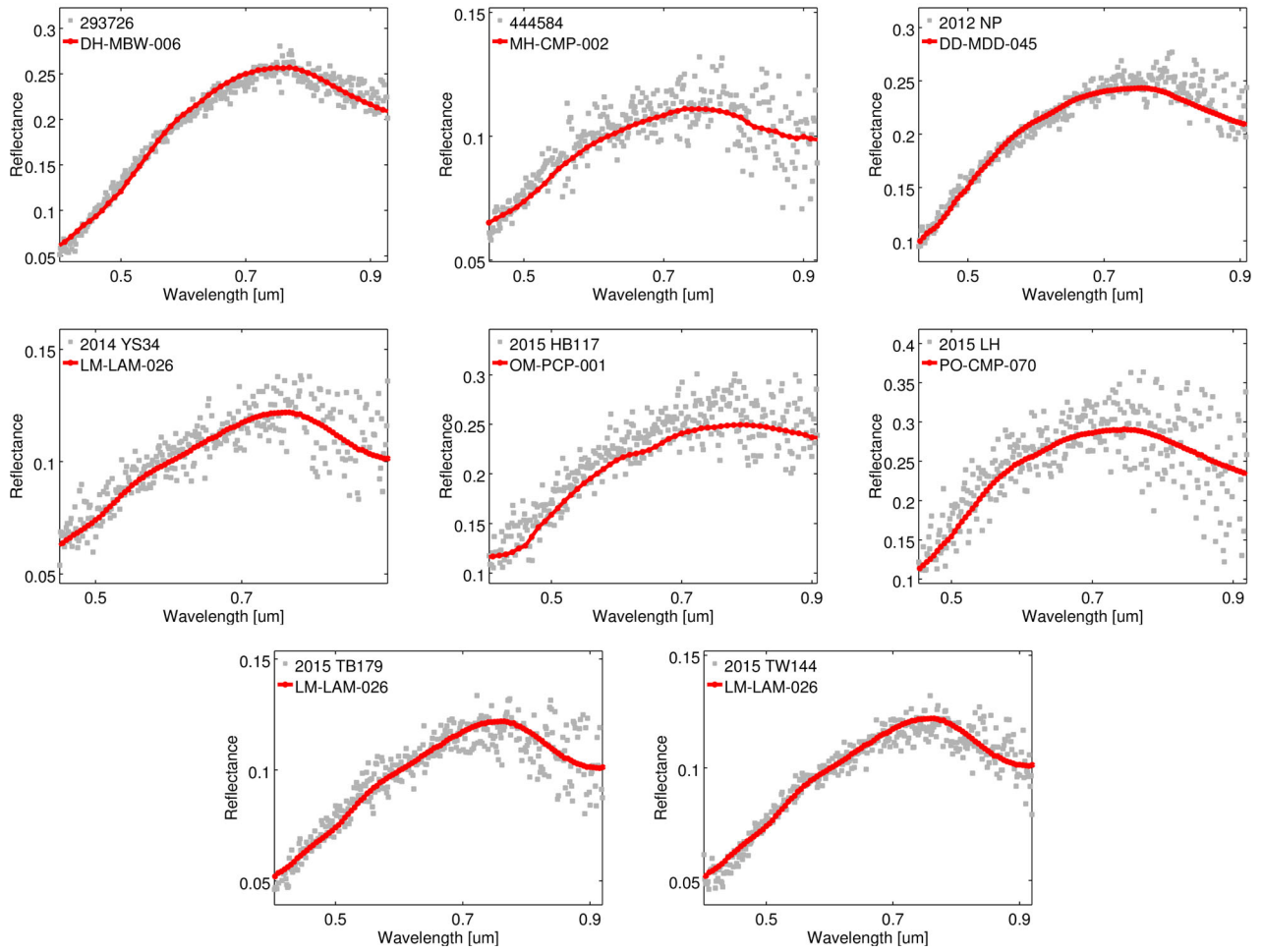
- Jedicke R., Larsen J., Spahr T., 2002, in Bottke W. F., Jr, Cellino A., Paolicchi P., Binzel R. P., eds, *Asteroids III, Observational Selection Effects in Asteroid Surveys*, p. 71
- Kohout T. et al., 2014, *Icarus*, 237, 75
- Kuroda D. et al., 2014, *PASJ*, 66, 51
- Mainzer A. et al., 2011, *ApJ*, 741, 90
- Menichella M., Paolicchi P., Farinella P., 1996, *Earth Moon Planets*, 72, 133
- Perna D. et al., 2018, *Planet. Space Sci.*, in press
- Polishook D., Jacobson S. A., Morbidelli A., Aharonson O., 2017, *Nature Astron.*, 1, 0179
- Popescu M., Birlan M., Nedelcu D. A., 2012, *A&A*, 544, A130
- Popescu M. et al., 2016, *A&A*, 591, A115
- Pravec P., Harris A. W., 2007, *Icarus*, 190, 250
- Reddy V., Nathues A., Gaffey M. J., Schaeff S., 2011, *Planet. Space Sci.*, 59, 772
- Sanchez J. A., Reddy V., Nathues A., Cloutis E. A., Mann P., Hiesinger H., 2012, *Icarus*, 220, 36
- Sanchez J. A. et al., 2014, *Icarus*, 228, 288
- Scheinberg A., Fu R. R., Elkins-Tanton L. T., Weiss B. P., 2015, in Patrick M., Francesca E. D., William F. B., eds, *Asteroid Differentiation: Melting and Large-Scale Structure*. University of Arizona Press, Tucson, p. 533
- Schulze H., Bischoff A., Palme H., Spettel B., Dreibus G., Otto J., 1994, *Meteoritics*, 29, 275
- Scott E. R. D., Keil K., Goldstein J. I., Asphaug E., Bottke W. F., Moskovitz N. A., 2015, in Patrick M., Francesca E. D., William F. B., eds, *Early Impact History and Dynamical Origin of Differentiated Meteorites and Asteroids*. University of Arizona Press, Tucson, 895, p. 573
- Singer R. B., 1981, *J. Geophys. Res.*, 86, 7967
- Sunshine J. M., Pieters C. M., 1998, *J. Geophys. Res.*, 103, 13675
- Sunshine J. M., Bus S. J., Corrigan C. M., McCoy T. J., Burbine T. H., 2007, *Meteorit. Planet. Sci.*, 42, 155
- Tholen D. J., 1984, *Asteroid Taxonomy from Cluster Analysis of Photometry*. University of Arizona Press, Tucson, AZ, p. 166
- Tholen D. J., Barucci M. A., 1989, in Binzel R. P., Gehrels T., Matthews M. S., eds, *Asteroids II*. University of Arizona Press, Tucson, AZ, p. 298
- Vaduvescu O. et al., 2018, *A&A*, 609, A105
- Veeder G. J., Matson D. L., Tedesco E. F., 1983, *Icarus*, 55, 177
- Whiteley R. J., Jr, 2001, PhD thesis, University of Hawai'i at Manoa

## APPENDIX A: SPECTRAL COMPARISON WITH DATA AVAILABLE IN RELAB DATA BASE

**Table A1.** Comparison with RELAB data base. The SampleID of the most relevant spectrum is given. Additional information including sample name, type, sub-type, and texture is given as they are provided by the RELAB data base. The spectra are shown in Fig. A1.

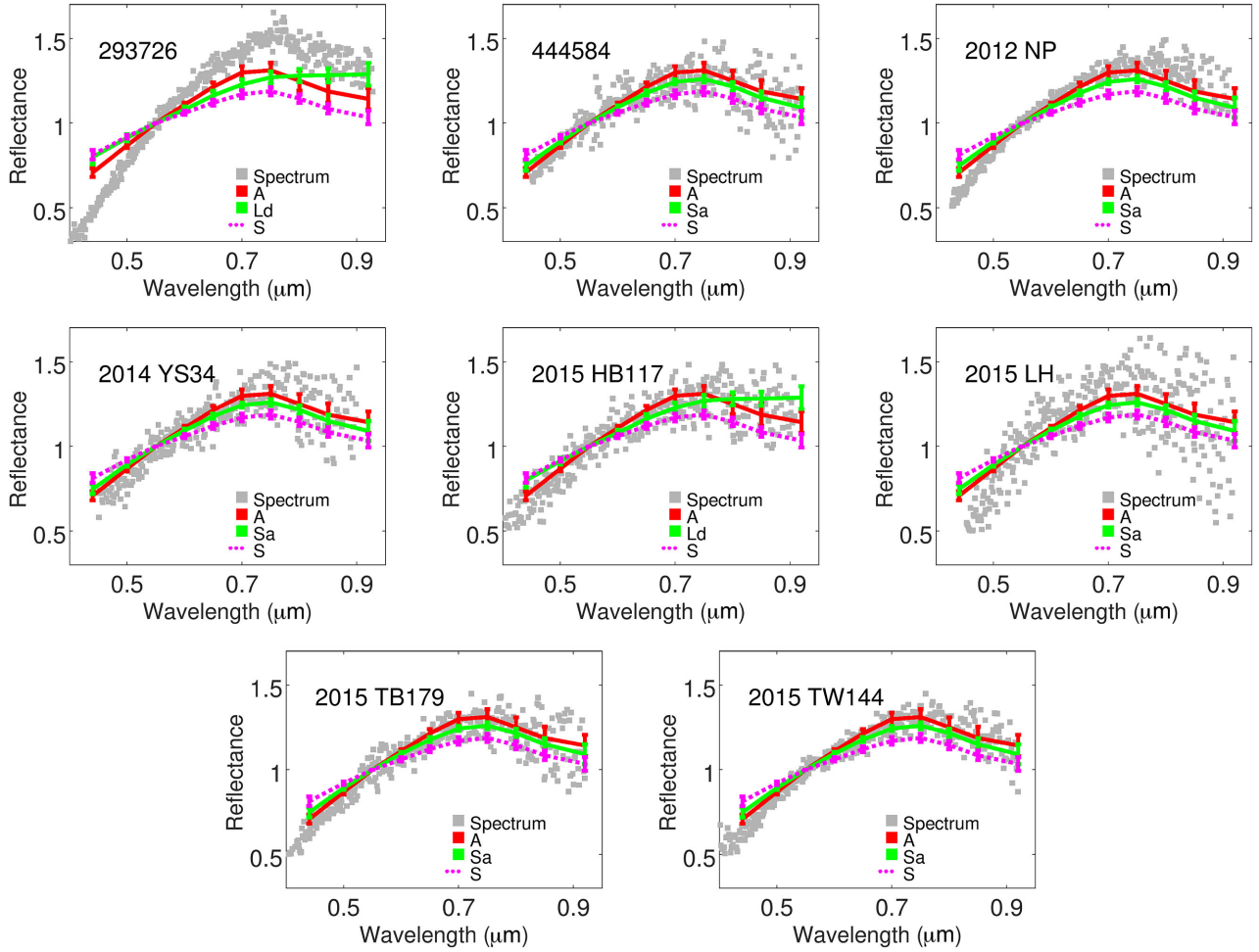
Asteroid	SampleID	Sample Name	General Type	Type	Sub Type	Texture
293726	DH-MBW-006	Fayalite 1–30 $\mu\text{m}$	Mineral	Silicate(Neso)	Olivine Fayalite	Particulate ( $\leq 30 \mu\text{m}$ )
444584	MH-CMP-002	Barratta	Rock	Ordinary Chondrite	L4	Slab
2012 NP	DD-MDD-045	Fo 10	Mineral	Silicate (Neso)	Olivine	Particulate ( $\leq 3 < 45 \mu\text{m}$ )
2014 YS34	LM-LAM-026	Dhajala	Rock	Ordinary Chondrite	H3-4 Olivine-Bronzite	Thin Section
2015 HB117	OM-PCP-001	01OP1c	Rock	Ultramafic	Dunite	Particulate (250 – 1000 $\mu\text{m}$ )
2015 LH	PO-CMP-070	St.Peter's Fayalite	Mineral	Silicate (Neso)	Olivine Fayalite	Particulate ( $\leq 3 < 45 \mu\text{m}$ )
2015 TB179	LM-LAM-026	Dhajala	Rock	Ordinary Chondrite	H3-4 Olivine-Bronzite	Thin Section
2015 TW144	LM-LAM-026	Dhajala	Rock	Ordinary Chondrite	H3-4 Olivine-Bronzite	Thin Section





**Figure A1.** Plot of the asteroid spectra versus the best match spectrum from the RELAB data base. The asteroid is identified by its designation and it is plotted in grey; the RELAB spectrum is shown in red and is referenced by its sample ID (see Table A1 for additional details).

## APPENDIX B: SPECTRAL COMPARISON WITH AVERAGE SPECTRA FROM BUS TAXONOMY



**Figure B1.** Matching of the observed spectra with taxonomic templates defined by Bus & Binzel (2002b). The visible spectra of (293726) 2007 RQ17, (444584) 2006 UK, 2012 NP, 2014 YS34, 2015 HB117, 2015 LH, 2015 TB179, 2015 TW144 are shown in grey. All spectra are normalized to  $0.55 \mu\text{m}$ . The first two best matches reported in Table 2 are shown. The S-type which is the most common type in the NEA population is plotted for comparison.

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.